

Ice Giants Revisited: Uranus and Neptune as Magma Ocean Worlds

Supplemental Information

April 3, 2026

Neptune Summary

Table 1: Neptune interior model summary. The model has two fixed inputs (Section 1) and two free parameters optimized by MCMC (Section 2). Section 3 compares model predictions to observations.

Fixed Inputs (set by Neptune)		
Parameter	Value	Source
Total mass M	1.024×10^{26} kg	Jacobson (2009) ^a
Equilibrium temperature T_{eq}	47 K	Voyager 2 ^b
Free Parameters (MCMC best fit)		
Parameter	Best-fit value	68% interval
Binodal surface pressure P_{binodal}	10.12 GPa	$9.20^{+0.67}_{-1.27}$ GPa
Bulk H ₂ mass fraction f_{H_2}	10.07 wt%	$9.73^{+0.25}_{-0.37}$ wt%
Model Predictions vs. Neptune Observations		
Observable	Model	Neptune observed
1-bar radius $R_{1\text{bar}}$	24,759 km	24,764 km ^c
Bulk density ρ	1610 kg m ⁻³	1638 kg m ^{-3c}
$J_2 \times 10^6$	3535.5	$3528.9 \pm 48.9^{\dagger,d}$
$J_4 \times 10^6$	-35.84	$-35.83 \pm 10.76^{\dagger,d}$
Norm. moment of inertia C/MR^2	0.2313	0.2315 ± 0.0046^e
Internal luminosity L_{int}	1.58×10^{15} W	$\sim 2-3 \times 10^{15}$ W ^f

[†] 1σ uncertainty used in the MCMC cost function (cal sigma mode): LOO-weighted in-sample calibration residual for J_2 (48.9×10^{-6} , 1.4%), floored at the observational sigma for J_4 (10.76×10^{-6} , 30%). J_2 and J_4 are consistency checks; all four observables are within 0.13σ at the maximum-likelihood point.

^a (Jacobson, 2009).

^b (Conrath et al., 1989) (Voyager 2 IRIS; equilibrium temperature derived from infrared measurements).

^c (Lindal, 1992) (Voyager 2 radio occultation; 1-bar equatorial radius 24 766 km adopted; bulk density computed from mass and volumetric mean radius).

^d J_2 and J_4 from (Wang et al., 2023), originally at $R_{\text{ref}} = 25,225$ km, renormalized to $R_{\text{eq}} = 24,766$ km following (Morf and Helled, 2025).

^e (Nettelmann et al., 2013); 2% uncertainty assumed.

^f (Pearl and Conrath, 1991).

MCMC Estimation for Two Free Parameters

Table 2: Posterior parameter estimates and maximum-likelihood solution.

Parameter	Posterior Median	68% Interval
Surface pressure P_{binodal} (GPa)	9.197	+0.666 / - 1.266
Bulk H ₂ fraction f_{H_2} (wt%)	9.726	+0.252 / - 0.373
<i>Maximum likelihood solution</i>		
P_{binodal} (GPa)		10.1240
f_{H_2} (wt%)		10.0659
$\ln p$		-0.0104
χ_{total}^2		0.0207
χ_{reduced}^2	0.0104	(d.o.f. = 2)

Predicted Observables vs. Neptune (Cost Function)

Table 3: Best-fit predicted observables compared to Neptune observations. All residuals are in units of the total 1σ uncertainty used in the cost function.

Observable	Predicted	Observed	Residual	Δ/σ
$J_2 \times 10^6$	3535.51	3528.91	+6.60	+0.13 σ
$J_4 \times 10^6$	-35.84	-35.83	-0.01	-0.00 σ
C/MR^2	0.2313	0.2315	-0.0002	-0.05 σ
$R_{1\text{bar}}$ (km)	24759.1	24764.0	-4.9	-0.02 σ

Table 4: Observational uncertainties used in the MCMC cost function (cal sigma mode). σ_{cal} for J_2 is the LOO-weighted in-sample residual (1.4%). σ_{cal} for J_4 equals the observational sigma (obs floor applied). σ_{total} is the value actually used in the cost function.

Quantity	σ_{obs}	σ_{cal}	σ_{total}	Basis
$J_2 \times 10^6$	4.14	48.94 (1.4%)	48.94	LOO-wt
$J_4 \times 10^6$	10.76	10.76 (30.0%)	10.76	obs floor
C/MR^2		± 0.0046 (assumed 2%)		—
$R_{1\text{bar}}$ (km)		± 247.6 (assumed 1%)		—

Predicted Bulk Planet Properties

Table 5: Bulk planet properties at the best-fit parameters.

Property	Value
Total mass, input Neptune value	1.024×10^{26} kg ($17.15 M_{\oplus}$)
Radius at 1-bar level (R_{\oplus})	3.886
Radius at 1-bar level (km)	24,759.2
Neptune observed bulk density	≈ 1638 kg m ⁻³
Normalized moment of inertia C/MR^2 (1-bar)	0.231

Predicted Thermal Properties

Table 6: Thermal properties of the best-fit model.

Property	Value
Equilibrium temperature T_{eq} , input Neptune value	47.0 K ^b
Radiation temperature T_{rad}	43.74 K
Internal luminosity L_{int}	1.580×10^{15} W
Neptune observed L_{int}	$\sim 2\text{--}3 \times 10^{15}$ W ^f
Central temperature	7271 K
Surface temperature (base of atmosphere)	2750 K
Mass-weighted temperature	1704 K
Temperature at radiative–convective boundary T_{rcb}	650 K

Predicted Core Structure

Table 7: Core (silicate + H₂–silicate mixture) properties.

Property	Value
Core mass (M_{\oplus})	16.196
Core radius (R_{\oplus})	3.042
Metal core radius (R_{\oplus})	0.000
Core volume (m ³)	3.049×10^{22}
Bulk core density (kg m ⁻³)	3172.3
H ₂ mass fraction in core	0.048
Thermal energy, core (J)	8.741×10^{32}
Gravitational potential energy, core (J)	-2.034×10^{34}

Predicted Atmosphere Structure

Table 8: Atmospheric envelope properties.

Property	Value
Atmosphere mass (kg)	5.700×10^{24}
Atmosphere mass fraction (M_{atm}/M_p)	5.565×10^{-2}
H ₂ gas mass fraction of planet	5.272%
H ₂ condensate mass fraction of planet	$2.342 \times 10^{-3}\%$
H ₂ mass fraction in atmosphere (gas)	0.9995
H ₂ mass fraction whole planet	0.1007
Surface pressure at base of atmosphere	1.012×10^5 bar (= 10.12 GPa)
Pressure at base of atmosphere (weight)	5.175×10^5 bar
Radiative-convective boundary height	4805 km
R_{rcb}/R_c	1.248
Pressure at rcb	201.4 bar
Optical depth at rcb	9.52×10^4
λ_{rcb}	120.12
Convection inhibited layer thickness	1.02×10^5 m
Boundary layer thickness, envelope	70.4 m
$R(1 \text{ bar})/R_c$	1.278
Atmosphere volume (1-bar, m ³)	3.309×10^{22}
Mean atm. density (1-bar, kg m ⁻³)	172.2
Planet volume fraction: core	0.474
Planet volume fraction: atmosphere	0.526

Predicted Total Energy Budget

Table 9: Integrated energy budget for the full planet.

Quantity	Value (J)
Thermal energy, core	$+8.741 \times 10^{32}$
Thermal energy, planet	$+9.249 \times 10^{32}$
Gravitational potential energy, core	-2.034×10^{34}
Gravitational potential energy, planet	-2.151×10^{34}
Total energy (planet)	-2.058×10^{34}

Uranus Summary

Table 10: Uranus interior model summary. The model has two fixed inputs (Section 1) and two free parameters optimized by MCMC (Section 2). Section 3 compares model predictions to observations.

Fixed Inputs (set by Uranus)		
Parameter	Value	Source
Total mass M	8.681×10^{25} kg	Jacobson & Park (2025) ^A
Equilibrium temperature T_{eq}	58 K	Pearl & Conrath (1990) ^B
Free Parameters (MCMC best fit)		
Parameter	Best-fit value	68% interval
Binodal surface pressure P_{binodal}	9.84 GPa	$9.23^{+0.61}_{-1.13}$ GPa
Bulk H ₂ mass fraction f_{H_2}	10.90 wt%	$10.47^{+0.35}_{-0.68}$ wt%
Model Predictions vs. Uranus Observations		
Observable	Model	Uranus observed
1-bar radius $R_{1\text{bar}}$	25,432 km	25,559 km ^C
$J_2 \times 10^6$	3507.30	$3509.29 \pm 58.9^{\dagger,D}$
$J_4 \times 10^6$	-35.56	$-35.52 \pm 0.466^{\dagger,D}$
Norm. moment of inertia C/MR^2	0.2217	0.2250 ± 0.0045^E
Bulk density ρ^{\ddagger}	1260 kg m ⁻³	≈ 1270 kg m ⁻³
Internal luminosity L_{int}^{\S}	5.29×10^{14} W	$\sim 6.4 \times 10^{14}$ W ^B

[†] Uncertainties used in the MCMC cost function (cal sigma mode). J_2 sigma (58.871×10^{-6} , 1.68%) is the LOO-weighted in-sample calibration residual. J_4 sigma (0.466×10^{-6} , 1.31%) equals the observational sigma (obs floor applied). For Uranus, $R_{\text{ref}} = R_{\text{eq}} = 25,559$ km (French et al. 2024 = Lindal et al. 1987), so published gravity values are already normalized to the 1-bar equatorial radius; no renormalization was applied.

[‡] Uranus observed bulk density ≈ 1270 kg m⁻³ uses the IAU volumetric mean radius.

[§] Internal luminosity based on $R_a/R_{a_c} = 10^{10}$.

^A (Jacobson and Park, 2025).

^B (Pearl and Conrath, 1990) (Voyager 2 IRIS; equilibrium temperature and internal luminosity for Uranus).

^C (Lindal et al., 1987) (Voyager 2 radio occultation; 1-bar equatorial radius 25 559 km).

^D (French et al., 2024); reported at $R_{\text{ref}} = R_{\text{eq}} = 25,559$ km — no renormalization required. Rotation period from (Desch et al., 1986).

^E (Nettelmann et al., 2013); see also (Podolak and Helled, 2012). 2% uncertainty assumed.

MCMC Estimation for Two Free Parameters

Table 11: Posterior parameter estimates and maximum-likelihood solution.

Parameter	Posterior Median	68% Interval
Surface pressure P_{binodal} (GPa)	9.228	+0.606 / - 1.125
Bulk H ₂ fraction f_{H_2} (wt%)	10.465	+0.346 / - 0.684
<i>Maximum likelihood solution</i>		
P_{binodal} (GPa)		9.8392
f_{H_2} (wt%)		10.8979
$\ln p$		-0.3948
χ_{total}^2		0.7897
χ_{reduced}^2	0.3948	(d.o.f. = 2)

Predicted Observables vs. Uranus (Cost Function)

Table 12: Best-fit predicted observables compared to Uranus observations. All residuals are in units of the total 1σ uncertainty used in the cost function.

Observable	Predicted	Observed	Residual	Δ/σ
$J_2 \times 10^6$	3507.29	3509.29	-2.00	-0.03 σ
$J_4 \times 10^6$	-35.56	-35.52	-0.03	-0.07 σ
C/MR^2	0.2217	0.2250	-0.0033	-0.73 σ
$R_{1\text{bar}}$ (km)	25431.4	25559.0	-127.6	-0.50 σ

Table 13: Observational uncertainties used in the MCMC cost function (cal sigma mode). σ_{cal} for J_2 is the LOO-weighted in-sample residual from the all-4-planet fit (Jupiter, Saturn, Uranus, Neptune), with ice giants carrying $\sim 99.8\%$ of the total weight. σ_{cal} for J_4 equals the observational sigma (obs floor applied). σ_{total} is the value used in the cost function.

Quantity	σ_{obs}	σ_{cal}	σ_{total}	Basis
$J_2 \times 10^6$	0.412	58.871 (1.68%)	58.871	LOO-wt
$J_4 \times 10^6$	0.466	0.466 (1.31%)	0.466	obs floor
C/MR^2		± 0.0045 (assumed 2%)		—
$R_{1\text{bar}}$ (km)		± 255.6 (assumed 1%)		—

Note: No J_2/J_4 renormalization was applied. For Uranus, $R_{\text{ref}} = R_{\text{eq}} = 25,559$ km (French et al. 2024 = Lindal et al. 1987), so the published values are already at the 1-bar equatorial radius. Contrast with Neptune, where $R_{\text{ref}} \neq R_{\text{eq}}$. J_2/J_4 forward model uses an empirical calibration with an all-4-planet fit: $J_2 = a q + b q C + c q^2$ (in-sample residuals $< 0.13\%$); $J_4 = \text{slope}_{\text{ice}} \times J_2$ (mean J_4/J_2 for Uranus + Neptune; residuals $< 1.4\%$).

Predicted Bulk Planet Properties

Table 14: Bulk planet properties at the best-fit parameters.

Property	Value
Total mass, input Uranus value	8.681×10^{25} kg ($14.536 M_{\oplus}$)
Radius at 1-bar level (R_{\oplus})	3.995
Radius at 1-bar level (km)	25,451.2
Uranus observed bulk density	≈ 1270 kg m ⁻³
Normalized moment of inertia C/MR^2 (1-bar)	0.223

Predicted Thermal Properties

Table 15: Thermal properties of the best-fit model.

Property	Value
Equilibrium temperature T_{eq} , input Uranus value	58.1 K ^B
Radiation temperature T_{rad}	54.07 K
Internal luminosity L_{int}	4.628×10^{14} W
Uranus observed L_{int}	$\sim 6.4 \times 10^{14}$ W ^B
Central temperature	7041.2 K
Surface temperature (base of atmosphere)	2826.7 K
Mass-weighted temperature	1829.1 K
Temperature at radiative–convective boundary T_{rcb}	710.6 K

Predicted Core Structure

Table 16: Core (silicate + H₂–silicate mixture) properties.

Property	Value
Core mass (M_{\oplus})	13.740
Core radius (R_{\oplus})	3.036
Metal core radius (R_{\oplus})	0.000
Core volume (m ³)	3.033×10^{22}
Bulk core density (kg m ⁻³)	2705.7
H ₂ mass fraction in core	0.057
Thermal energy, core (J)	$+7.277 \times 10^{32}$
Gravitational potential energy, core (J)	-1.461×10^{34}

Predicted Atmosphere Structure

Table 17: Atmospheric envelope properties.

Property	Value
Atmosphere mass (kg)	4.758×10^{24}
Atmosphere mass fraction (M_{atm}/M_p)	5.481×10^{-2}
H ₂ gas mass fraction of planet	5.196%
H ₂ condensate mass fraction of planet	$1.152 \times 10^{-3}\%$
H ₂ mass fraction in atmosphere (gas)	0.9983
H ₂ mass fraction whole planet	0.1087
Surface pressure at base of atmosphere	9.434×10^4 bar (= 9.434 GPa)
Pressure at base of atmosphere (weight)	4.010×10^5 bar
Radiative–convective boundary height	5164 km
R_{rcb}/R_c	1.267
Pressure at rcb	655.8 bar
Optical depth at rcb	7.92×10^5
λ_{rcb}	91.97
Convection inhibited layer thickness	1.45×10^5 m
Boundary layer thickness, envelope	393.7 m
$R(1 \text{ bar})/R_c$	1.316
Atmosphere volume (1-bar, m ³)	3.873×10^{22}
Mean atm. density (1-bar, kg m ⁻³)	122.8
Planet volume fraction: core	0.434

Predicted Total Energy Budget

Table 18: Integrated energy budget for the full planet.

Quantity	Value (J)
Thermal energy, core	$+7.145 \times 10^{32}$
Thermal energy, planet	$+7.745 \times 10^{32}$
Gravitational potential energy, core	-1.465×10^{34}
Gravitational potential energy, planet	-1.580×10^{34}
Total energy (planet)	-1.502×10^{34}

References

- B. J. Conrath, F. M. Flasar, R. Hanel, V. Kunde, W. Maguire, J. Pearl, J. Pirraglia, R. Samuelson, P. Gierasch, and A. Weir. Infrared observations of the Neptunian system. *Science*, 246:1454–1459, 1989. doi:[10.1126/science.246.4936.1454](https://doi.org/10.1126/science.246.4936.1454). URL <https://doi.org/10.1126/science.246.4936.1454>.

- M. D. Desch, J. E. P. Connerney, and M. L. Kaiser. The rotation period of Uranus. *Nature*, 322: 42–43, 1986. doi:[10.1038/322042a0](https://doi.org/10.1038/322042a0). URL <https://doi.org/10.1038/322042a0>.
- R. G. French, M. M. Hedman, P. D. Nicholson, P.-Y. Longaretti, and C. A. McGhee-French. The Uranus system from occultation observations (1977–2006): Rings, pole direction, gravity field, and masses of Cressida, Cordelia, and Ophelia. *Icarus*, 411:115957, 2024. doi:[10.1016/j.icarus.2024.115957](https://doi.org/10.1016/j.icarus.2024.115957). URL <https://doi.org/10.1016/j.icarus.2024.115957>.
- R. A. Jacobson. The orbits of the Neptune trojans and the mass of Neptune. *The Astronomical Journal*, 137:4322–4329, 2009. doi:[10.1088/0004-6256/137/5/4322](https://doi.org/10.1088/0004-6256/137/5/4322). URL <https://doi.org/10.1088/0004-6256/137/5/4322>.
- R. A. Jacobson and R. S. Park. The masses of Uranus and its major satellites from spacecraft and earth-based observations. *The Astronomical Journal*, 169:65, 2025. doi:[10.3847/1538-3881/ad9b4d](https://doi.org/10.3847/1538-3881/ad9b4d). URL <https://doi.org/10.3847/1538-3881/ad9b4d>.
- G. F. Lindal. The atmosphere of Neptune: an analysis of radio occultation data acquired with Voyager 2. *The Astronomical Journal*, 103:967–982, 1992. doi:[10.1086/116119](https://doi.org/10.1086/116119). URL <https://doi.org/10.1086/116119>.
- G. F. Lindal, J. R. Lyons, D. N. Sweetnam, V. R. Eshleman, D. P. Hinson, and G. L. Tyler. The atmosphere of Uranus: Results of radio occultation measurements with Voyager 2. *Journal of Geophysical Research*, 92:14987–15001, 1987. doi:[10.1029/JA092iA13p14987](https://doi.org/10.1029/JA092iA13p14987). URL <https://doi.org/10.1029/JA092iA13p14987>.
- L. Morf and R. Helled. Icy or rocky? Convective or stable? New interior models of Uranus and Neptune. *Astronomy & Astrophysics*, 704:A183, 2025. doi:[10.1051/0004-6361/202556911](https://doi.org/10.1051/0004-6361/202556911). URL <https://doi.org/10.1051/0004-6361/202556911>.
- N. Nettelmann, J. J. Fortney, K. Moore, and C. Mankovich. An assessment of giant planet interior models with an application to Saturn’s interior. *Planetary and Space Science*, 77:143–151, 2013. doi:[10.1016/j.pss.2012.06.019](https://doi.org/10.1016/j.pss.2012.06.019). URL <https://doi.org/10.1016/j.pss.2012.06.019>.
- J. C. Pearl and B. J. Conrath. The albedo, effective temperature, and energy balance of Uranus, as determined from Voyager IRIS data. *Journal of Geophysical Research*, 95:18921–18930, 1990. doi:[10.1029/JA095iA12p18921](https://doi.org/10.1029/JA095iA12p18921). URL <https://doi.org/10.1029/JA095iA12p18921>.
- J. C. Pearl and B. J. Conrath. The albedo, effective temperature, and energy balance of Neptune, as determined from Voyager IRIS data. *Journal of Geophysical Research*, 96:18921–18930, 1991. doi:[10.1029/91JA01087](https://doi.org/10.1029/91JA01087). URL <https://doi.org/10.1029/91JA01087>.
- M. Podolak and R. Helled. What do we really know about Uranus and Neptune? *The Astrophysical Journal Letters*, 759:L32, 2012. doi:[10.1088/2041-8205/759/2/L32](https://doi.org/10.1088/2041-8205/759/2/L32). URL <https://doi.org/10.1088/2041-8205/759/2/L32>.
- B. Wang, J. Yan, W. Gao, Y. Yuan, S. Sun, M. Ye, and J.-P. Barriot. The Neptunian gravity estimated from the motion of Triton based on astrometric observations. *Astronomy & Astrophysics*, 671:A70, 2023. doi:[10.1051/0004-6361/202244537](https://doi.org/10.1051/0004-6361/202244537). URL <https://doi.org/10.1051/0004-6361/202244537>.